

THE CHALLENGE OF SUB-KEPLERIAN ROTATION FOR DISK WINDS

FRANK H. SHU,¹ SUSANA LIZANO,² DANIELE GALLI,³ MIKE J. CAI,⁴ AND SUBHANJOY MOHANTY⁵*Received 2007 November 20; accepted 2008 June 11; published 2008 July 16*

ABSTRACT

Strong magnetization makes the disks surrounding young stellar objects rotate at rates that are too sub-Keplerian to enable the thermal launching of disk winds from their surfaces unless the rate of gas diffusion across field lines is dynamically fast. This underappreciated implication of disk magnetization poses a considerable challenge for disk-wind theory.

Subject headings: accretion, accretion disks — ISM: jets and outflows — stars: pre-main-sequence

1. INTRODUCTION

The pioneering study by Blandford & Payne (1982; see also Chan & Henriksen 1980) opened the door to a physical understanding of the highly collimated jets that emanate from magnetized disks accreting onto central gravitating masses. Pudritz & Norman (1983) made the first application to the disks in star formation, while Shibata & Uchida (1985) performed the first numerical simulations. Collectively, the disk-wind community has contributed to much of what we know about the magnetohydrodynamic (MHD) processes of jet launching, acceleration, and collimation.

Invoking strong magnetization to aid in the magnetohydrodynamics of the gas above the disk has serious consequences for the gas contained within the disk. Wardle & Königl (1993, hereafter WK93), Ferreira & Pelletier (1995, FP95), Ferreira (1997, F97), Casse & Ferreira (2000, CF00), and Salmeron et al. (2007, SKW07) performed the primary studies of some of the effects. The result has been a shrinkage of the viable parameter space to launch disk winds. In fact, unless the gas diffuses across magnetic field lines at dynamical rates, the sub-Keplerian rotation of the surfaces of realistic disks in young stellar objects (YSOs) may make magnetocentrifugally driven disk winds a severe challenge if their surfaces are cold.

In § 2 we give the basic equations from the analysis of Shu et al. (2007, S07). In § 3 we use this analysis to demonstrate that fast disk winds that are more than lightly loaded will not arise if they are thermally launched. In § 4 we review how WK93, FP95, F97, CF00, and SKW07 formally escape the conclusions of § 3 by assuming that the diffusion of gas across magnetic field lines inside the disk occurs on dynamical time-scales, and we discuss the astrophysical difficulties engendered by such a point of view. In § 5 we summarize how X-wind theory surmounts the same difficulties.

2. BASIC EQUATIONS

Use a cylindrical coordinate system (ϖ, φ, z) centered on a star with mass M_* , and consider centrifugal balance in a thin

disk rotating at radius ϖ with an angular velocity $\Omega(\varpi)$ that is a density-weighted average over z :

$$-\Sigma\varpi\Omega^2 = \frac{B_z B_\varpi^+}{2\pi} - \frac{GM_*\Sigma}{\varpi^2}, \quad (2.1)$$

where G is the universal gravitational constant, Σ is the surface density, and B_z and B_ϖ^+ are, respectively, the midplane and upper surface values of the vertical and radial components of the magnetic field. From Ampère's law in cgs units, $cB_\varpi^+/2\pi$ equals the vertically integrated current density J_φ , where c is the speed of light. Equation (2.1) then assumes that the Lorentz force per unit area, $J_\varphi B_z/c$, associated with magnetic tension, provides the only opposition, apart from inertia, to the attraction of stellar gravity. We have ignored the self-gravity of the disk, which would additionally hinder the launching of magnetocentrifugal winds. When integrating the divergence of the Maxwell stress tensor over the disk thickness, we have also ignored terms like $-(1/8\pi)\partial(B_\varphi^2)/\partial\varpi$ as being a factor of the disk aspect ratio $\ll 1$ smaller than the retained term $(1/4\pi)\partial(B_z B_\varpi)/\partial z$ for a spatially thin disk. Retention of the effects of magnetic and gas pressure, which generally decrease with increasing radius ϖ , would only worsen the problem of sub-Keplerianity detailed below.

With \bar{f} denoting the z -averaged fraction of Keplerian rotation at ϖ , $\Omega \equiv \bar{f}\Omega_K$ where $\Omega_K \equiv (GM_*/\varpi^3)^{1/2}$, we obtain (S07, Appendix C)

$$1 - \bar{f}^2 = \frac{\varpi^2 B_z B_\varpi^+}{2\pi GM_* \Sigma}. \quad (2.2)$$

Magnetic effects also affect the vertical structure of the disk. We define a characteristic thermal velocity a such that $a^2 \equiv 2z_0 P(0)/\Sigma$, where $P(0)$ is the midplane gas pressure. Integrating the equation of vertical hydrostatic equilibrium in z gives (S07, Appendix C)

$$\frac{(B_\varpi^+)^2}{2\pi} = \Sigma \left(\frac{2a^2}{A\varpi} - \frac{GM_* A}{\varpi^2} \right), \quad (2.3)$$

where we define A and A_0 as, respectively, the actual and thermal aspect ratios of the disk,

$$A \equiv \frac{z_0}{\varpi}, \quad A_0 \equiv \left(\frac{2a^2 \varpi}{GM_*} \right)^{1/2}. \quad (2.4)$$

In the above, z_0 is a vertical height of the disk defined so that the z -integrated value of $z\rho$ from the midplane to the surface

¹ Department of Physics, University of California, San Diego, La Jolla, CA 92093; fshu@physics.ucsd.edu.

² CRYA, Universidad Nacional Autónoma de México, 58089 Morelia, Mexico.

³ INAF-Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, Firenze I-50125, Italy.

⁴ Academia Sinica, Institute of Astronomy and Astrophysics, Taiwan.

⁵ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138.

of zero gas-pressure equals $(z_0/2)(\Sigma/2)$. Note that $A_0 = \sqrt{2}(a/\varpi\Omega_K) \ll 1$ in a thin disk where $\varpi\Omega_K$ is the Keplerian speed at radius ϖ .

WK93, FP95, F97, CF00, and SKW07 never employ the relation (2.2) to estimate \bar{f} . Instead, they allow departures from hydrostatic equilibrium above the disk that involve advective accelerations in the radial and vertical directions that are $O(A_0)$ when measured in a nondimensional sense. Section 4 demonstrates that such departures occur only if there is a transonic diffusion of gas across magnetic field lines that include a strong component of B_φ not present in the analysis of S07.

Denoting $B_\varpi^+/B_z = I_l$ with $B_z \propto \varpi^{-(1+l)}$ in self-similar models, and substituting equation (2.3) into equation (2.2), we get

$$1 - \bar{f}^2 = \frac{A_0}{I_l} \left(\frac{A_0}{A} - \frac{A}{A_0} \right). \quad (2.5)$$

The presence of a radial field B_ϖ that increases from zero in the midplane to B_ϖ^+ at the surface compresses the disk and makes A smaller than A_0 . In the process, $1 - \bar{f}^2$ becomes greater than zero.

To make contact with disk-wind theory, let us follow FP95 in defining a “magnetization,”

$$\mu \equiv \frac{B_z^2}{4\pi P(0)} = \frac{B_z^2 z_0}{2\pi \Sigma a^2}, \quad (2.6)$$

which also equals a quantity that WK93 call α_0^2 . Dividing equation (2.3) by equation (2.6) now yields $I_l^2 = (2/\mu)(1 - A^2/A_0^2)$, where we have replaced B_ϖ^+/B_z by I_l . Using this relation to eliminate A/A_0 from equation (2.5) now gives

$$1 - \bar{f}^2 = \frac{A_0}{I_l} \left[\frac{1}{C_l(\mu)} - C_l(\mu) \right], \quad (2.7)$$

where the positive quantity $C_l(\mu) = A/A_0$ is $C_l(\mu) \equiv (1 - I_l^2 \mu/2)^{1/2} \leq 1$.

The maximum value of μ allowed mechanically in equation (2.7) comes by setting $\bar{f} = 0$. The aspect ratio then assumes the minimum value:

$$A_{\text{mech}} = \frac{I_l}{2} [\sqrt{1 + 4(A_0/I_l)^2} - 1]. \quad (2.8)$$

From $C_l(\mu_{\text{mech}}) = A_{\text{mech}}/A_0$, we obtain μ_{mech} :

$$\mu_{\text{mech}} = \frac{1}{A_0^2} [\sqrt{1 + 4(A_0/I_l)^2} - 1]. \quad (2.9)$$

For $\mu > \mu_{\text{mech}}$, magnetic forces are too strong for centrifugal equilibrium to be possible. For $\mu = \mu_{\text{mech}}$ in thin disks where $A_0/I_l \ll 1$, the roots (2.8) and (2.9) take the approximate forms, $A_{\text{mech}} \approx A_0^2/I_l$, $\mu_{\text{mech}} \approx 2/I_l^2$.

In the model of Blandford & Payne (1982) and in cases where the wind is relatively lightly loaded (e.g., F97), $l = 1/4$ or nearly so, i.e., $B_z \propto \varpi^{-5/4}$. The magnetic configuration just above the disk is then well approximated by a vacuum field, where the analysis of S07 yields $I_l = 1.43$. The corresponding inclination angle of the field from the vertical is $i = 55^\circ$. With $I_l = 1.43$, $\mu_{\text{mech}} \approx 0.978$. For μ not closely approaching μ_{mech} , say, $\mu \approx 0.731$ where the magnetic compression $C(\mu) \approx 0.5$, the right-hand side of equation (2.7) is

$O(A_0) \ll 1$. The analogous feature in numerical calculations, combined with the ameliorating effects of rapid diffusion, led disk-wind theorists to assert that the departure from Keplerian rotation remains small enough for other effects to overcome the deficit.

3. THERMAL LAUNCHING

Thermal pressure cannot do the requisite job in the absence of fast (turbulent) diffusion (see § 4). Published models have the deficit of f from unity at the compressed surface $z = z_0$ both larger than average (e.g., Fig. 5 of FP95) and smaller than average (e.g., Fig. 1 of SKW07). As a fiducial case, we assume the disk's surface to have the z -averaged f obtained from equation (2.7). To have a significant pressure exhaust of gas onto open field lines, we then require the thermal speed squared a_s^2 at the surface of the disk to be (S07, eq. [31])

$$a_s^2 = \frac{1}{4} (1 - \bar{f}^2) \frac{GM_*}{\varpi} = \frac{1}{2I_l A_0} \left[\frac{1}{C_l(\mu)} - C_l(\mu) \right] a^2, \quad (3.1)$$

where we applied equations (2.4) and (2.7) in the last step.

Without external heating, the surface temperature of the disk, $\sim 2^{-1/4}$ times the effective temperature, will be lower than its characteristic interior value, i.e., $a_s^2 = \Theta_s a^2$ where Θ_s is a fraction $\lesssim 1$. Equation (3.1) then yields a quadratic equation for $C_l(\mu_{\text{therm}}) = A_{\text{therm}}/A_0$, whose solution is

$$A_{\text{therm}} = A_0 [\sqrt{1 + (I_l \Theta_s A_0)^2} - (I_l \Theta_s A_0)]. \quad (3.2)$$

From $C_l(\mu_{\text{therm}}) = A_{\text{therm}}/A_0$, we obtain μ_{therm} :

$$\mu_{\text{therm}} = 4 \frac{\Theta_s A_0}{I_l} [\sqrt{1 + (I_l \Theta_s A_0)^2} - (I_l \Theta_s A_0)]. \quad (3.3)$$

For thin disks where $I_l \Theta_s A_0 \ll 1$, equations (3.2) and (3.3) give $A_{\text{therm}} \approx A_0$, $\mu_{\text{therm}} \approx 4\Theta_s A_0/I_l$. The right-hand side of the second relation is, at most, several percent in the inner tenth of an AU in YSO disks. The magnetization is then so weak that the associated departure from Keplerian rotation from equation (2.7), $1 - \bar{f}^2 \approx 2\Theta_s A_0^2$, would be similar to that implied by the radial gradients of the thermal gas pressure, i.e., very small. But weak magnetization, $\mu \leq \mu_{\text{therm}}$, with good field loading, does not provide a long enough lever arm to launch fast disk winds even if $\bar{f} = 1$ (see § 3.5 of CF00).

If the disk magnetization μ is chosen greater than μ_{therm} , but less than μ_{mech} , the actual departure $1 - \bar{f}^2 \sim O(A_0)$ exceeds the depth of the dimensionless effective potential $1 - \bar{f}^2 \sim O(A_0^2)$ that the surface temperature can offset. Because a_s^2 may be less than 1% of GM_*/ϖ in the inner disk, even $\bar{f} = 0.99$ may constitute a substantial barrier by these standards. As argued earlier, no room exists for a compromise μ between μ_{therm} and μ_{mech} that will drive disk winds that both load and fling if one depends on thermal pressure to launch.

4. DIFFUSIVE LOADING

WK93, FP95, F97, CF00, and SKW07 differ with the analysis of §§ 2 and 3 because they do not invoke thermal pressure to launch disk winds. Instead, they overcome local deficits in f by diffusive loading onto magnetic field lines coupled to magnetocentrifugal acceleration.

For an imperfectly conducting, lightly ionized, gas, suffi-

ciently collisional to act as a single-component fluid of velocity \mathbf{v} , the induction equation reads

$$\frac{\partial \mathbf{B}}{\partial t} + \nabla \times (\mathbf{B} \times \mathbf{v}) = -\nabla \times [\eta \cdot (\nabla \times \mathbf{B})], \quad (4.1)$$

where η is a generalized tensor resistivity associated with Ohm's law (Norman & Heyvaerts 1985, see especially eq. [22]). In a frame that corotates locally with the magnetic field, so that $\partial \mathbf{B} / \partial t = 0$, we may “uncurl” equation (4.1) and obtain the balance of advection and diffusion, $\mathbf{B} \times \mathbf{u} = -\eta \cdot (\nabla \times \mathbf{B})$, where \mathbf{u} is the fluid velocity relative to the rotating field lines.

Keeping only the dominant terms in a thin disk, we follow WK93 and SKW07 in decomposing the advection-diffusion balance into its φ and ϖ components:

$$B_z u_\varpi - B_\varpi u_z = \eta_2 \frac{\partial B_\varphi}{\partial z} - \eta_1 \frac{\partial B_\varpi}{\partial z}, \quad (4.2)$$

$$B_\varphi u_z - B_z (f - f_B) \sqrt{\frac{GM_*}{\varpi}} = \eta_1 \frac{\partial B_\varphi}{\partial z} - \eta_2 \frac{\partial B_\varpi}{\partial z}, \quad (4.3)$$

where the dimensionless rotation of the field f_B is a function of z at fixed ϖ (in such an Eulerian description, we are not following a given field line). For WK93, η_1 is associated with the coefficient of ambipolar diffusion, and η_2 with the inverse of the Hall conductivity. In F97 and CF00, η_1 is replaced by ν_m in equation (4.2) and by ν'_m in equation (4.3), with the other, nondiagonal, coefficient η_2 set to zero.

In the disk proper (where u_z is small) of both Figure 1 of SKW07 and Figure 5 of FP95, the terms on the right-hand sides of equations (4.2) and (4.3) contribute, respectively, negatively to the radial drift and positively to the tangential drift relative to the field lines. These relative motions reverse as the gas climbs vertically into the disk wind. Eventually, the right-hand sides vanish as ideal MHD is approached (η_1 and $\eta_2 \rightarrow 0$), with \mathbf{u} becoming parallel to \mathbf{B} in the corotating frame. The physical interpretation of the result is as follows.

A field line rotates at a dimensionless angular velocity f_B that is intermediate between the rotation rate f of the gas near the midplane layers of the disk and the gas in the upper atmosphere that blows into a disk wind. This rate is fixed in steady state so that angular momentum balance is reached following a given field line (a condition not examined in this Letter). In the WK93 description, the neutral gas near the midplane would rotate at the Kepler speed except for collisions with ions and electrons that are tied to the more slowly rotating magnetic field lines. The collisions therefore make the neutral gas lose angular momentum, which then cause the neutrals to drift radially inward relative to the radially stationary field lines and the ions and electrons tied to them.

Freezing to field lines in the upper atmosphere of the disk yields the familiar effect of a “sliding bead on a rotating rigid wire” (Chan & Henriksen 1980) that accelerates the gas through the slow MHD transition as it gains angular momentum at the expense of the matter in the deeper layers of the disk. Thus, the gas that participates in the disk wind ends up rotating at super-Keplerian speeds. A similar phenomenon occurs in the FP95 description, except that the role of collisions between

neutrals and ions/electrons is replaced by a mixing of fluid parcels created by MHD turbulence.

When an outwardly bending field line rotates uniformly along its length, there will always be a point where centrifugal effects on an electrically conducting test particle placed there counterbalances an opposing gravitational field. In Figure 5 of FP95 and Figure 1 of SKW07, this point occurs at $\sim 2v_T/\Omega_K$, where v_T is the thermal velocity in the midplane. For a vertically isothermal disk (the most optimistic case), where $a = (2/\sqrt{\pi})v_T$, this point is $\sim 2.5z_0$ since $z_0 \sim 0.5A_0\varpi \sim 0.8v_T/\Omega_K$ for typical parameters used to launch disk winds. The height z_0 is defined by where the condition of vertical hydrostatic equilibrium would have created a surface. As in any atmosphere, the gas density does not really drop abruptly to zero at z_0 , but unless diffusion works very efficiently, there cannot be much gas at $2.5z_0$.

Trouble on this point enters when numerical solutions are adjusted to give large rates of diffusion. In order of magnitude, let us estimate $\partial B_\varpi / \partial z \sim B_\varpi^+ / z_0$ and $\partial B_\varphi / \partial z \sim B_\varphi^+ / z_0$ in the disk proper. In the more successful cases considered by FP95, the quantities $\nu_m \sim z_0 a \sim \nu'_m$. With $B_\varpi^+ > 0$ and $-B_\varphi^+ > 0$ comparable in magnitude to $B_z > 0$, equations (4.2) and (4.3), with u_z negligible in the disk proper, now yield the estimates,

$$u_\varpi \sim -a, \quad f - f_B \sim \frac{a}{\sqrt{GM_*/\varpi}} \sim A_0. \quad (4.4)$$

The first deduction, that the inward drift speed in the bulk of the disk approaches sonic values, is verified, for example, by Figure 4 of FP95 where $u_\varpi \approx -v_T$. The second deduction, that the dimensionless rotation rate of the gas f can exceed that of the field f_B by $O(A_0)$, would produce favorable conditions for the launching of disk winds if the field rotates slower than the gas in the disk proper, but still within fraction $O(A_0)$ of Keplerian.

In the work of FP95, the diffusion coefficients η_1 and η_2 are chosen large enough so that, unlike S07, the magnetic field is not well coupled to the matter in the deep interior of the disk. With $\eta_1 \sim z_0 a$, the diffusive timescale to straighten poloidal field lines that are bent on a scale z_0 is $\sim z_0^2/\eta_1 \sim z_0/a$, which is of order the inverse Kepler rate Ω_K^{-1} (or shorter, if the disk is severely magnetically compressed). The dynamical slippage of the matter past the field, both in the ϖ and φ directions, is a serious weakness of the models.

Rapid diffusion can carry gas to the slow MHD transition point beyond which magnetocentrifugal fling works in the outward direction to produce an outflowing wind. Such a strategy generates, through equations (4.2) and (4.4), radial drift speeds inside the disk that are transonic $\sim -v_T$. Consider what this implies if a disk wind were responsible for the primary accretion mechanism of the disk all the way out to an outer disk radius ~ 300 AU in a low-mass YSO. The midplane temperature at 300 AU is unlikely to be much less than, say, 20 K, which corresponds to an isothermal sound speed in cosmic molecular gas of $v_T = 0.278 \text{ km s}^{-1}$. The whole disk would then shrink onto the star in $300 \text{ AU} / (0.278 \text{ km s}^{-1}) \sim 5000 \text{ yr}$, which is far too short a timescale to be realistic.

If global magnetorotational instability is a viable alternative for YSO accretion disks, the analysis of S07 and Shu et al. (2008) implies that the turbulent resistivity η_1 is smaller than estimated by FP95 or F97 by an additional factor $\sim \mu A \ll 1$. If we continue to assume that $-B_\varphi^+$ and B_ϖ^+ are comparable to

B_z , with the former holding only if disk winds can be launched, then equations (4.2) and (4.3) imply

$$u_{\infty} \sim -\mu Aa, \quad f - f_B \sim \mu AA_0 \lesssim A_0^2. \quad (4.5)$$

The second relation above states that diffusion can now make up deficits of f_B only by amounts $\lesssim A_0^2$, where f_B itself is given by f to $O(A_0^2)$ in such circumstances. In turn, f departs from Keplerian rotation typically by $O(A_0)$ for strongly magnetized disks according to equation (2.7). A low (S07; see also Lubow et al. 1994) rather than a high (FP95 or F97) rate of field diffusion yields equation (3.1) as, after all, the correct approximate criterion for wind loading and launching.

The numerical choices made by WK93 are more conservative than those of FP95, and the net results for u_{∞} and $f - f_B$ are intermediate between the high- and low-diffusion scenarios described above (see Appendix of WK93). Specifically, in the numerical case depicted in Figure 1 of SKW07, where $\alpha_0 \equiv \sqrt{\mu} = 0.95$ and $a_s^2 = v_T^2 = 0.01 GM_*/\varpi$, the dimensionless loading is $\kappa = 3.2 \times 10^{-4}$ and the square of the lever arm is $\lambda = 395$ (these are β^{-1} and J_w , respectively, in the notation of X-wind theory), corresponding indeed to a disk wind that can fling, but is very lightly loaded. Such winds do not describe observed YSO jets well. The order-of-magnitude changes (exponential sensitivity) in quantities like κ and λ when η changes by a factor of a few gives powerful testimony to the crucial role that magnetic diffusivity plays in current theories of the launching of disk winds.

5. DISCUSSION AND CONCLUSION

Disk winds have been invoked as a solution to the angular momentum problem in accretion disks (e.g., Pelletier & Pudritz 1992). The challenge of sub-Keplerian rotation facing magnetocentrifugally driven, cold, disk winds for much, if not most,

of its radial span suggests that internal mechanisms other than the back reaction to disk winds must account for the primary accretion mechanism of such disks. Examples of such mechanisms include spiral density waves for high-mass disks and the magnetorotational instability for low-mass disks.

Disruption of the disk at an inner edge by a stellar magnetosphere will drive an X-wind (Cai et al. 2008 and references therein) if the resistivity in the disk is appreciably smaller than the viscosity (for numerical simulations, see Romanova et al. 2008). How does X-wind theory escape the conundrum posed by equation (2.1)? Near the inner edges of disks, the X field geometry involves a swing of outwardly bending field lines (right half of X) to inwardly bending field lines (left half of X) as accretion disks are truncated by funnel flows through their interactions with a stellar magnetosphere. The swing of $B_{\infty}^+ B_z$ from positive values through zero (vertical field lines that are “dead” to outflow or inflow) automatically promotes a transition from sub-Keplerian to Keplerian rotation. Indeed, when the density decreases sharply inward on a fractional radial scale $\sim A_0$, the no-longer-ignorable, extra, radial push of gas pressure toward the inner edge speeds up disk rotation and makes the footpoints of some of the outwardly bending field lines exterior to the dead zone sufficiently close to Keplerian rotation as to enable the thermal launching of X-winds (see the discussion of Shu et al. 2008).

We thank Sylvie Cabrit, Jonathan Ferreira, Al Glassgold, Arie Königl, Raquel Salmeron, and Mark Wardle for comments that improved the presentation of the Letter. F. S. acknowledges support from the Physics Department of UCSD, S. L. from CONACyT 48901 and PAPIIT-UNAM IN106107, D. G. from INAF-OAA and EC grant MRTN-CT-2006-035890, M. C. from a grant from the NSC to TIARA, and S. M. from a Spitzer Fellowship.

REFERENCES

- Blandford, R. D., & Payne, D. G. 1982, MNRAS, 199, 883
 Cai, M., Lin, H. H., Shang, H., & Shu, F. H. 2008, ApJ, 672, 489
 Casse, F., & Ferreira, J. 2000, A&A, 353, 1115 (CF00)
 Chan, K. L., & Henriksen, R. N. 1980, ApJ, 241, 534
 Ferreira, J. 1997, A&A, 319, 340 (F97)
 Ferreira, J., & Pelletier, G. 1995, A&A, 295, 807 (FP95)
 Lubow, S. H., Papaloizou, J., & Pringle, J. E. 1994, MNRAS, 267, 235
 Norman, C., & Heyvaerts, J. 1985, A&A, 147, 247
 Pelletier, G., & Pudritz, R. E. 1992, ApJ, 394, 117
 Pudritz, R. E., & Norman, C. A. 1983, ApJ, 274, 677
 Romanova, M. M., Long, M., Kulkarni, A. K., Kurosawa, R., Ustyugova, G. V., Koldoba, A. K., & Lovelace, R. V. E. 2008, in IAU Symp. 243, Star-Disk Interaction in Young Stars, ed. I. Appenzeller & J. Bouvier (Dordrecht: Kluwer), in press
 Salmeron, R., Königl, A., & Wardle, M. 2007, MNRAS, 375, 177
 Shibata, K., & Uchida, Y. 1985, PASJ, 37, 31
 Shu, F. H., Galli, D., Lizano, S., & Cai, M. 2008, in IAU Symp. 243, Star-Disk Interaction in Young Stars, ed. I. Appenzeller & J. Bouvier (Dordrecht: Kluwer), in press
 Shu, F. H., Galli, D., Lizano, S., Glassgold, A. E., & Diamond, P. 2007, ApJ, 665, 535 (S07)
 Wardle, M., & Königl, A. 1993, ApJ, 410, 218 (WK93)